Cosmological stochastic Higgs stabilization

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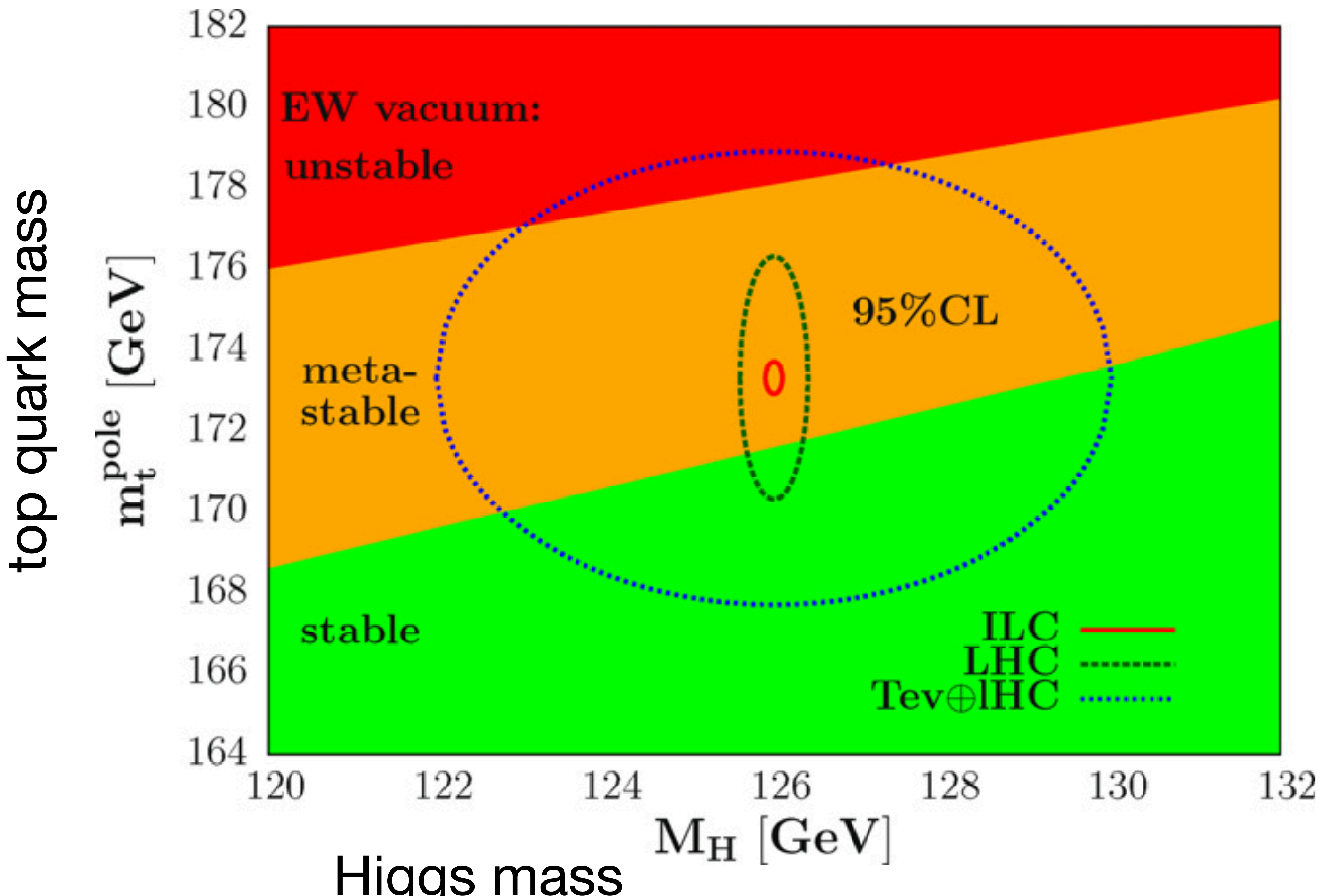
arXiv:1705.11178

COSMO-17, Paris

Higgs & top masses, vacuum stability

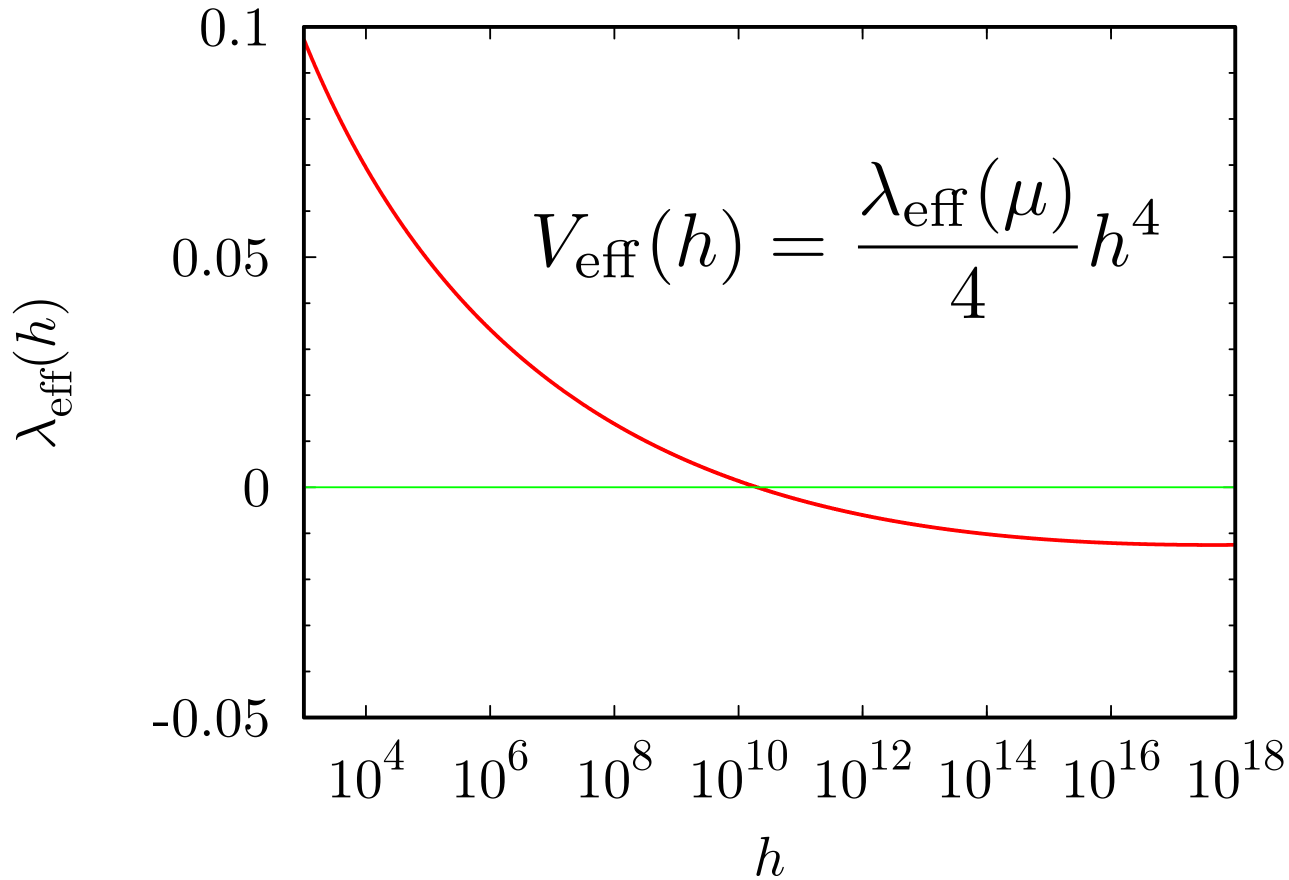


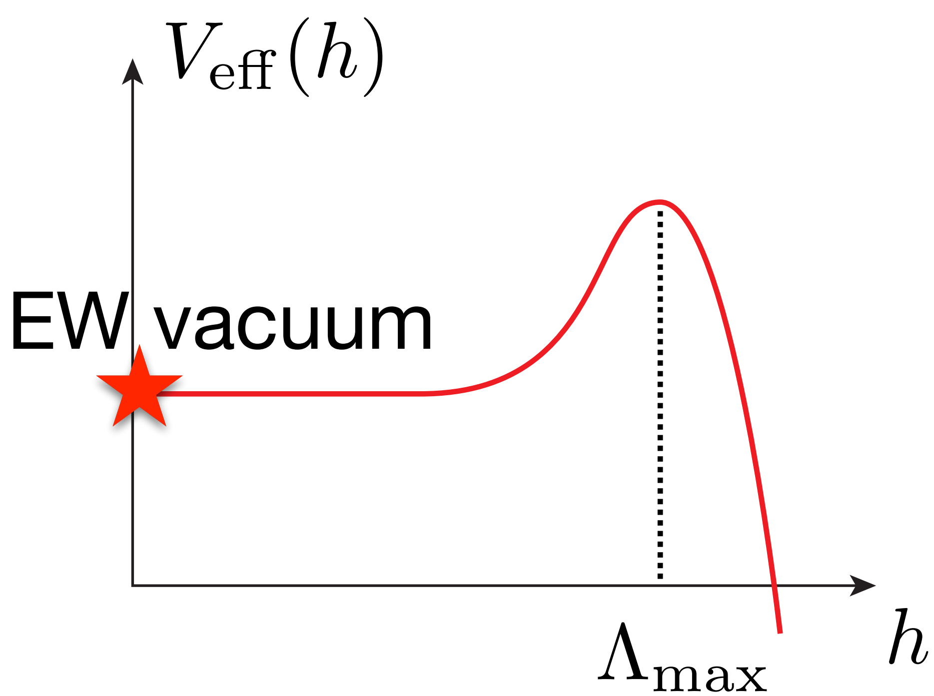
Alekhin, Djouadi, Moch, 1207.0980

Higgs & top masses, Higgs effective potential

$$m_h = 125.09 \pm 0.21 \pm 0.11 \text{ GeV}$$

$$m_t = 173.34 \pm 0.76 \pm 0.3 \text{ GeV}$$





$$\Lambda_{\rm max} \sim 10^{10-11} \text{ GeV}$$

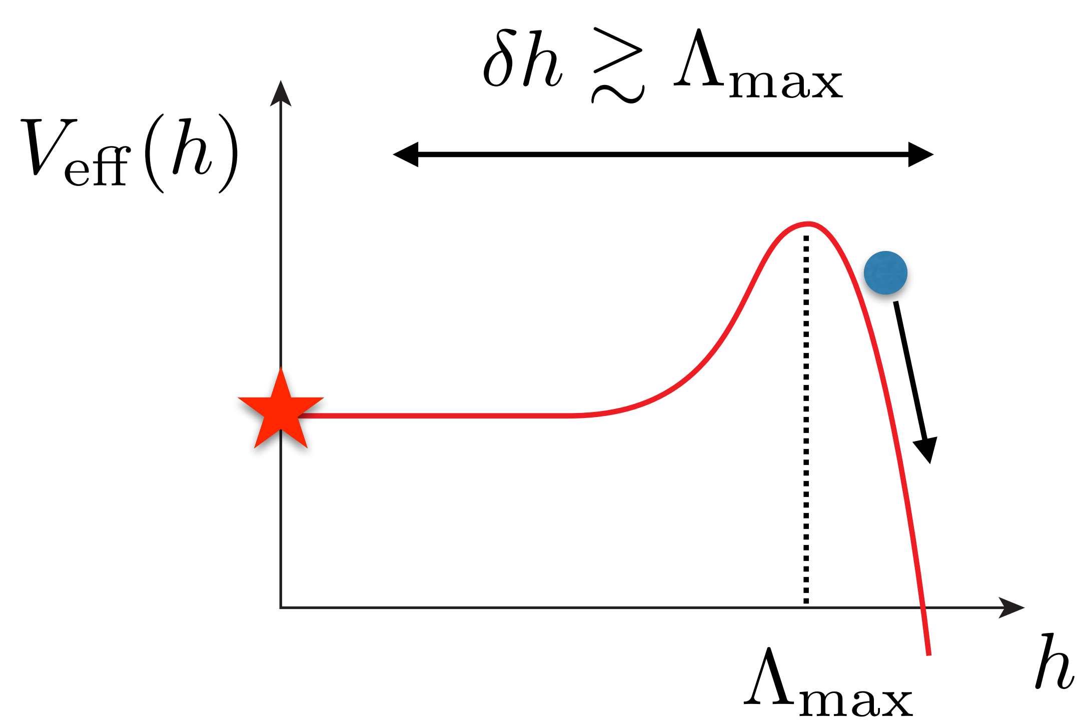
(if SM is valid up to some high energy scales)

Inflation and quantum fluctuation

Hubble parameter during inflation

$$H_{\rm inf} \lesssim A_{\mathcal{R}} \pi \sqrt{\frac{r}{2}} M_P \approx 6.73 \times 10^{13}\ \text{GeV} \left(\frac{r}{0.07}\right)^{1/2}$$

Quantum fluctuation of Higgs field : $\delta h \sim \dfrac{H_{\rm inf}}{2\pi}$



EW vacuum stabilization during inflation

- Low-scale inflation : $H_{\rm inf} \lesssim 10^{10-11}{\rm GeV}$

- Inflaton — Higgs coupling : $\Delta\mathcal{L} = -\frac{1}{2}\lambda_{\phi h}\phi^2 h^2$

Lebedev, Westphal, 1210.6987

- Non-minimal coupling with gravity : $\Delta\mathcal{L} = -\frac{1}{2}\xi R h^2$

Espinosa, Giudice, Riotto, 0710.2484

$\lambda_{\phi h}\phi^2,\ \xi R \gtrsim H_{\rm inf}^2$ —> Higgs is stabilized during inflation

However, parametric resonance / tachyonic instability after inflation destabilizes the EW vacuum. (—> Ema's talk)

Herranen, Markkanen, Nurmi, Rajantie, 1407.3141
Ema, Mukaida, Nakayama, 1602.00483

Stochastic dynamics — Fokker-Planck equation

$$\frac{\partial P(h,N)}{\partial N} = \frac{\partial}{\partial h}\left[\frac{V'(h)}{3H^2}P(h,N) + \frac{H^2}{8\pi^2}\frac{\partial P(h,N)}{\partial h}\right]$$

(IR-mode)

initial condition : $P(h,0) = \delta(h)$

....

classical motion : $\Delta h_{\rm cl} = \dot{h}\Delta t = -\frac{V'_{\rm eff}(h)}{3H^2}$

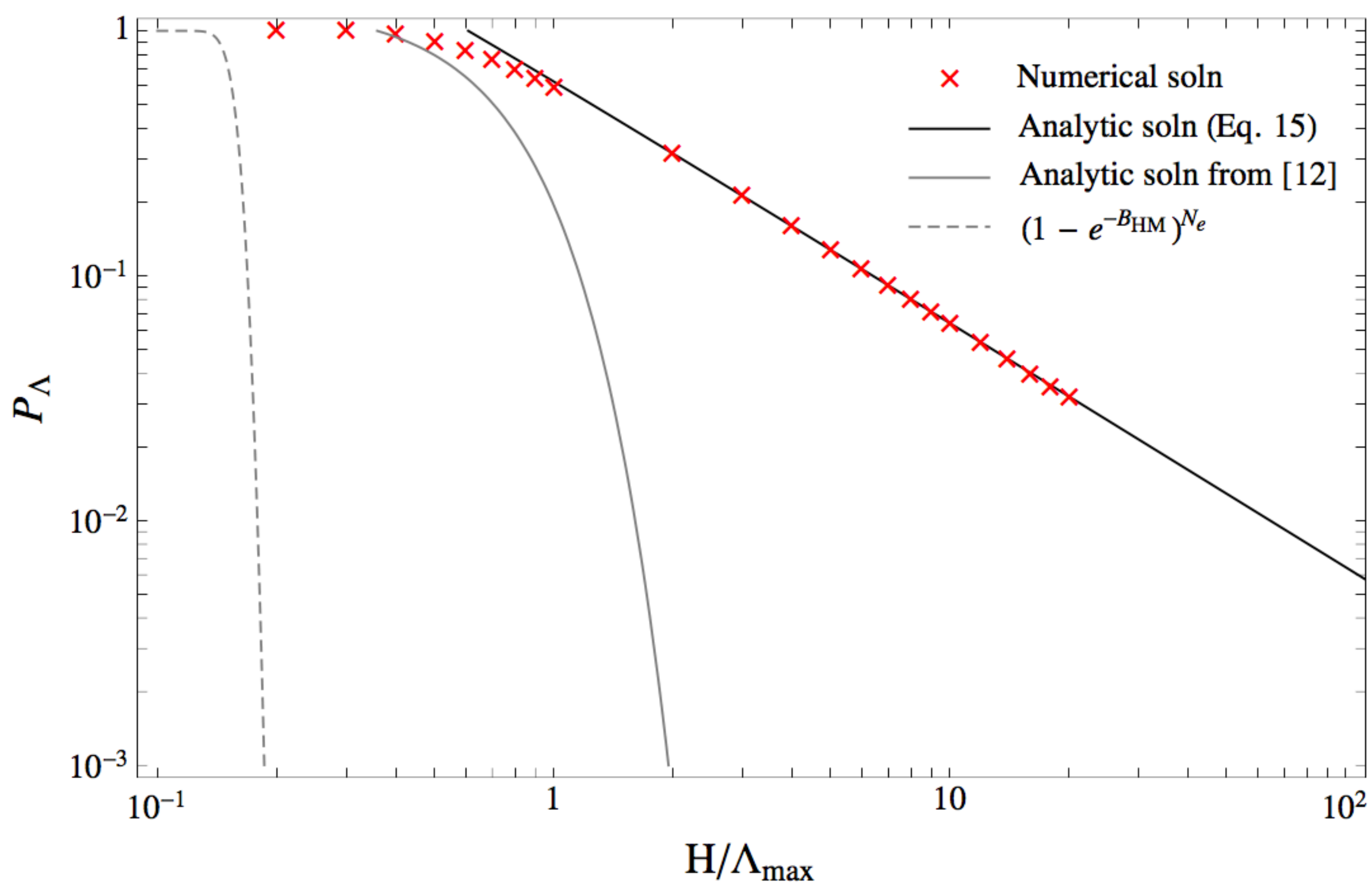
quantum fluctuation : $\Delta h_{\rm qm} = \frac{H}{2\pi}$

↓ $P(h,N) = 0$ for $\Delta h_{\rm qm} < \Delta h_{\rm cl}$

boundary condition : $|V'_{\rm eff}(\pm\Lambda_c)| \equiv \frac{3H^2}{2\pi}$

Survival probability

$$P_\Lambda = \int_{-\Lambda_{\max}}^{\Lambda_{\max}} dh P(h, N)$$



Hook, Kearney, Shakya, Zurek, 1404.5953

Multi-field stochastic dynamics during inflation

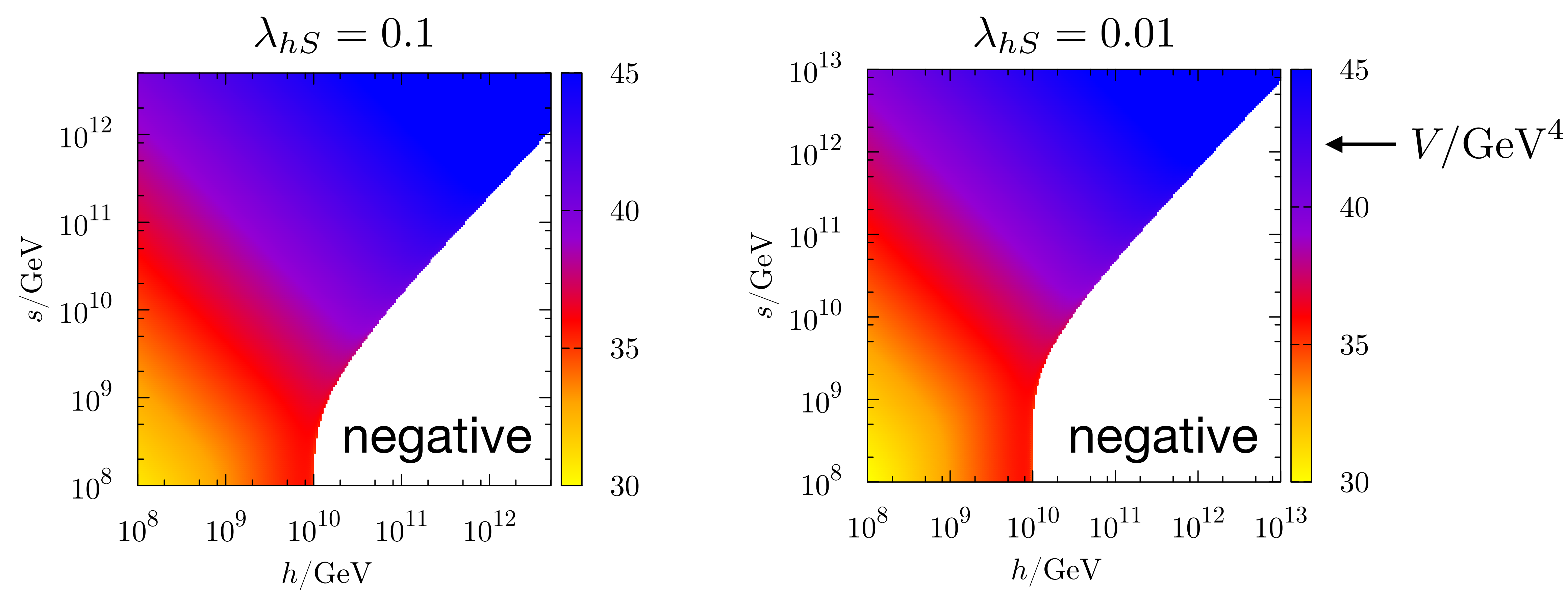
J.O. Gong, NK, 1705.11178

Model : Higgs (h) & singlet <u>spectator</u> field (S)

Potential : $$V(h,S) = \frac{1}{4}\lambda_{\rm eff}(\mu)h^4 + \frac{1}{2}m_S^2 S^2 + \frac{1}{2}\lambda_{hS}h^2S^2 + \frac{1}{4}\lambda_S S^4$$

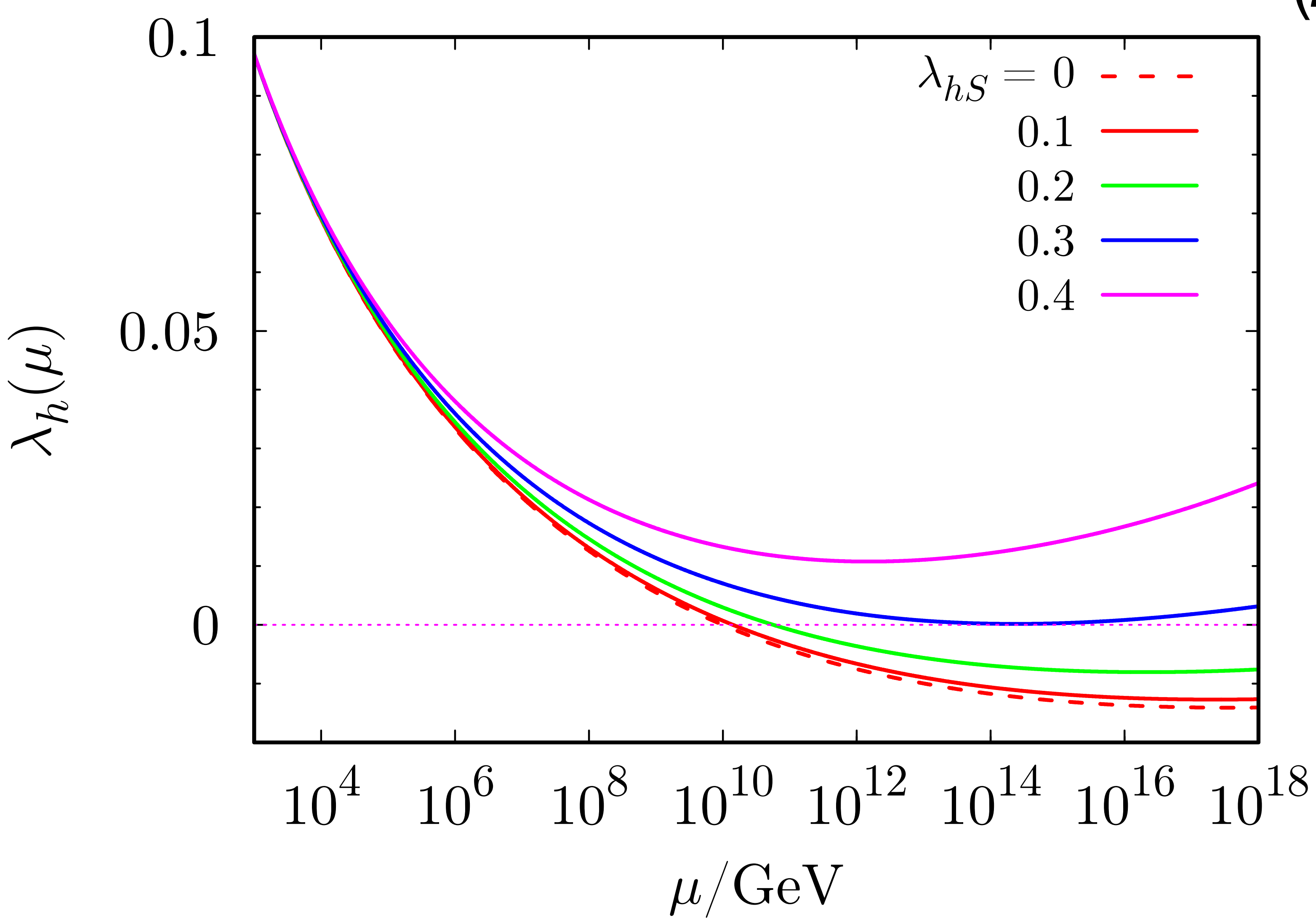
with $$\mu = \sqrt{h^2 + H_{\rm inf}^2}$$

Potential



Higgs self coupling in the presence of h-S coupling

(2-loop order)



$\lambda_{hS} \gtrsim 0.3$ —> EW vacuum is absolutely stable

Otherwise, the EW vacuum is still metastable…

Multi-field Fokker-Planck equation (IR-mode)

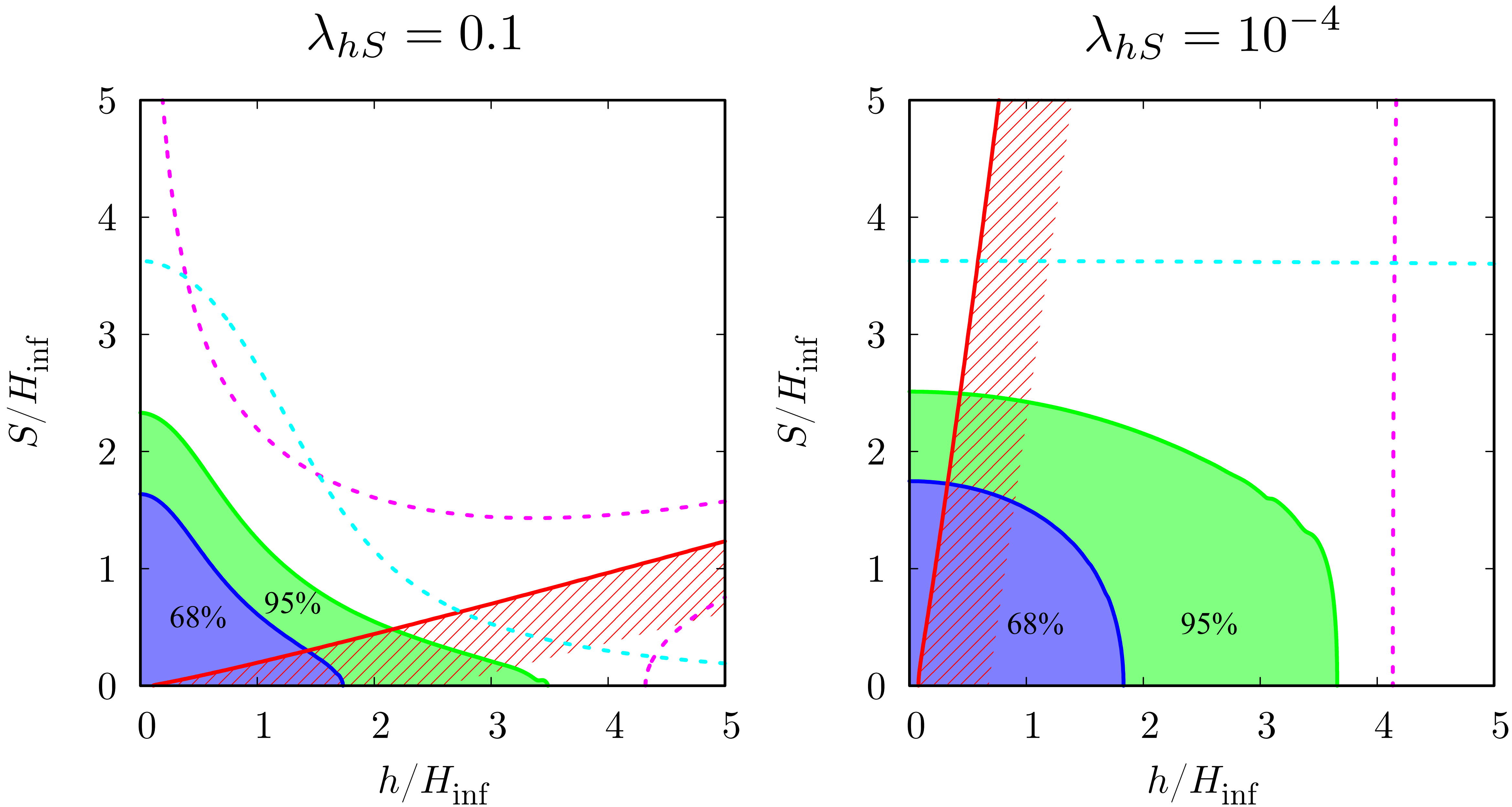
$$\frac{\partial P(h,S,N)}{\partial N} = \sum_{\phi=h,S} \frac{\partial}{\partial \phi} \left[\frac{\partial V(h,S)}{\partial \phi} \frac{P(h,S,N)}{3H^2} + \frac{H^2}{8\pi^2} \frac{\partial P(h,S,N)}{\partial \phi} \right]$$

with initial distribution : $P(h,S,0) = \delta(h)\delta(S)$

+ boundary condition : $P(\pm\Lambda_h, S, N) = P(h, \pm\Lambda_S, N) = 0$

with $\left|\frac{\partial V}{\partial h}\right|_{|h|=\Lambda_h} = \left|\frac{\partial V}{\partial S}\right|_{|S|=\Lambda_S} = \frac{3H_{\rm inf}^3}{2\pi}$

Equal probability contours



$$H_{\rm inf} = 10^{11}\ \text{GeV},\ N = 60,\ \lambda_S = 0.01$$

Analytic (approximate) solution

$$\frac{\partial P(h,S,N)}{\partial N} = \sum_{\phi=h,S} \frac{\partial}{\partial \phi}\left[\frac{\partial V(h,S)}{\partial \phi}\frac{P(h,S,N)}{3H^2} + \frac{H^2}{8\pi^2}\frac{\partial P(h,S,N)}{\partial \phi}\right]$$

separation of variables : $P(h,S,N) = \Psi(N)\Phi(h,S)$

↓

2nd term ↙

$\Psi(N) \propto e^{-\alpha N}$ & $\left(\frac{\partial^2}{\partial h^2} + \frac{\partial^2}{\partial S^2}\right)\Phi(h,S) = -\frac{8\pi^2}{H_{\rm inf}^2}\alpha\Phi(h,S)$

$$P(h,S,N) = \frac{1}{\Lambda_h \Lambda_S}\sum_{n,m\geq 0} e^{-\alpha_{nm}N}\cos\left[\pi\left(n+\frac{1}{2}\right)\frac{h}{\Lambda_h}\right]\left[\pi\left(m+\frac{1}{2}\right)\frac{S}{\Lambda_S}\right]$$

with $\alpha_{nm} = \left(n+\frac{1}{2}\right)^2\frac{H_{\rm inf}^2}{8\Lambda_h^2} + \left(m+\frac{1}{2}\right)^2\frac{H_{\rm inf}^2}{8\Lambda_S^2}$

survival probability : $P_\Lambda = \int_{-\Lambda_S}^{\Lambda_S} dS \int_{-\Lambda_{\max}}^{\Lambda_{\max}} dh P(h,S)$

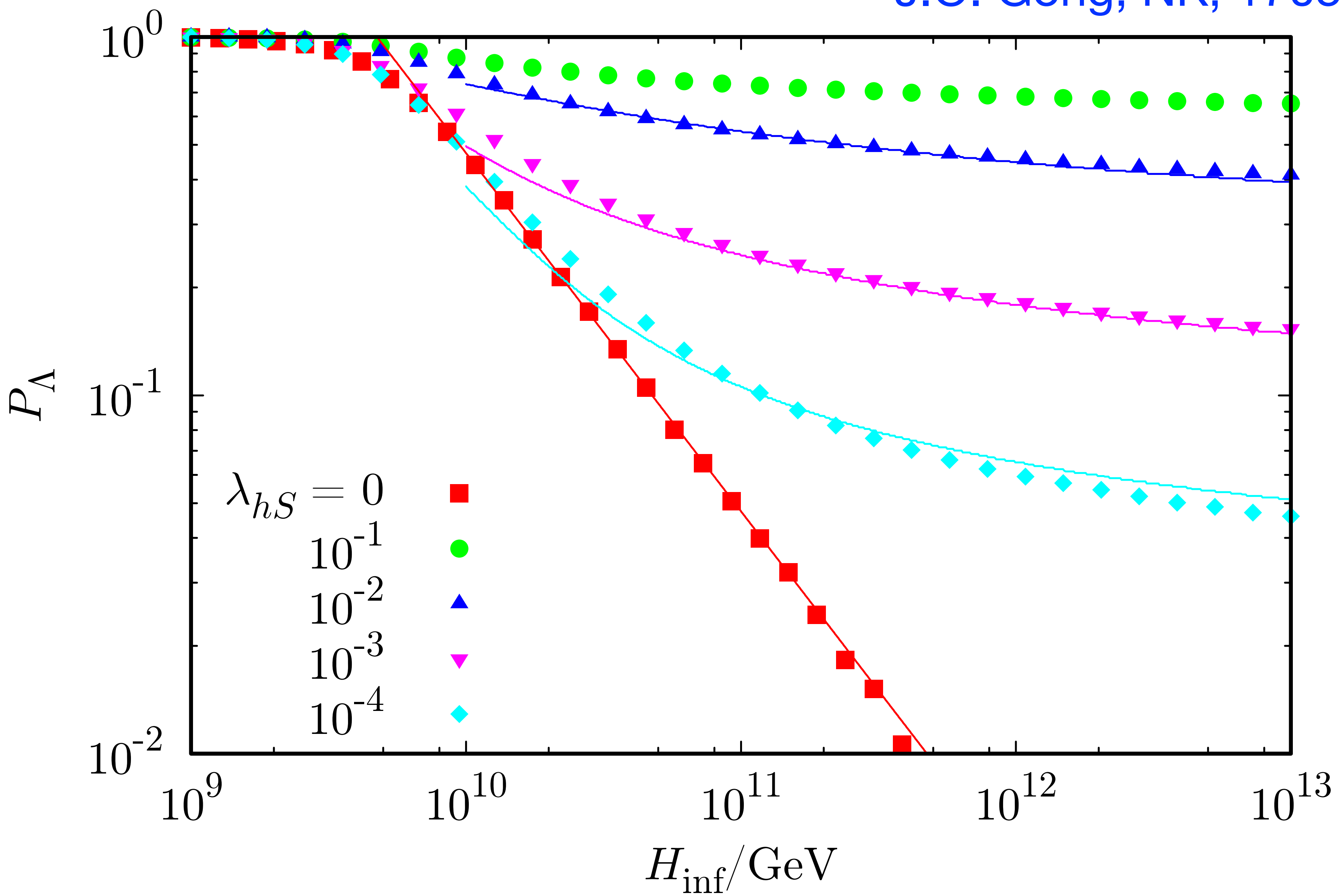
$$\Lambda_{\max} = \sqrt{\frac{\lambda_{hS}}{-\partial\lambda_{\rm eff}/4\partial\ln h - \lambda_{\rm eff}}} S \approx 10\sqrt{\lambda_{hS}} S$$

$$P_\Lambda = \frac{4p}{\pi^2}\sum_{n,m\geq 0} e^{-q\alpha_{nm}N}\frac{-\frac{m+1/2}{n+1/2}\Lambda_h^2 \sin\left[\beta\pi\left(n+\frac{1}{2}\right)\frac{\Lambda_S}{\Lambda_h}\right] + \beta\Lambda_s\Lambda_S}{\beta^2(n+1/2)^2\Lambda_S^2 - (m+1/2)^2\Lambda_h^2}$$

$$\beta = \Lambda_{\max}/S, \quad p, q \, : \text{ numerical fudge factors}$$

Survival probability just after inflation

J.O. Gong, NK, 1705.11178



$$N = 60,\ \lambda_S = 0.01$$

analytic curve -> $p = 2,\ q = 5$

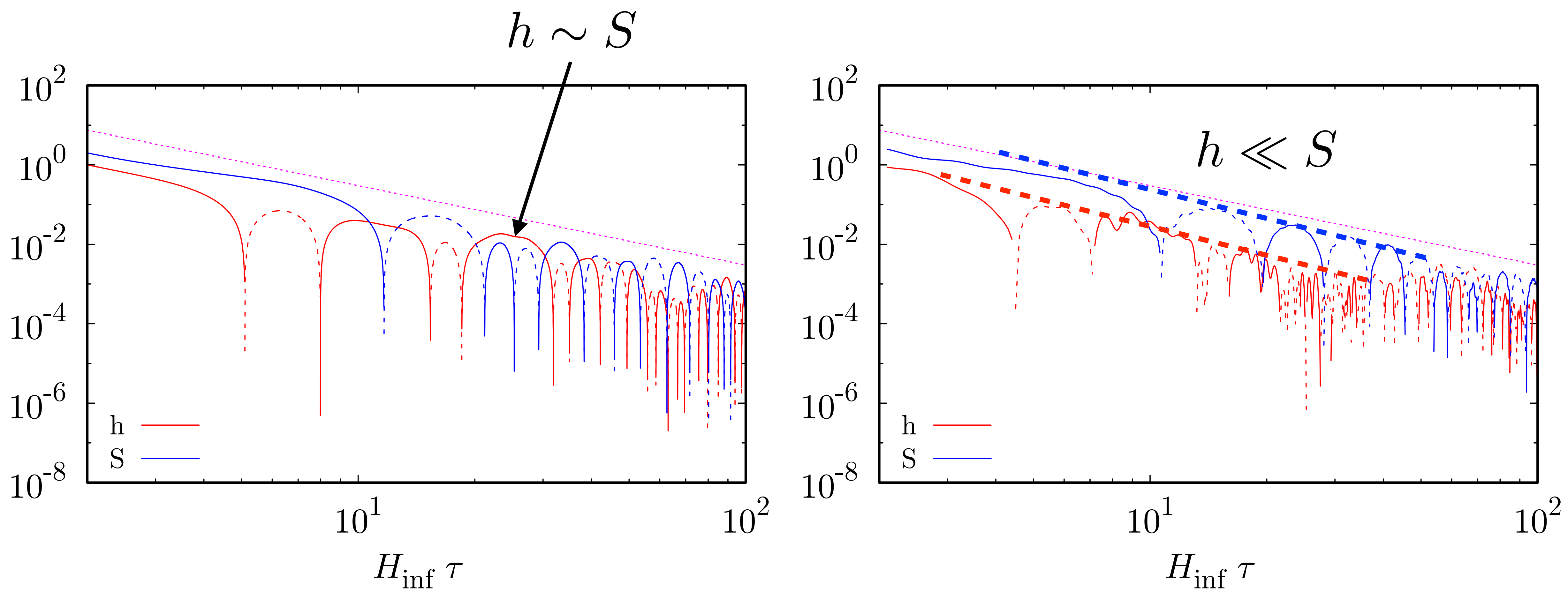
Scalar field dynamics after inflation

$$\ddot{h} + 3H\dot{h} - \frac{\nabla^2 h}{a^2} + \left(\lambda_{\rm eff} + \frac{1}{4}\frac{\partial \lambda_{\rm eff}}{\partial \ln h}\right) h^3 + \lambda_{hS} S^2 h = 0$$

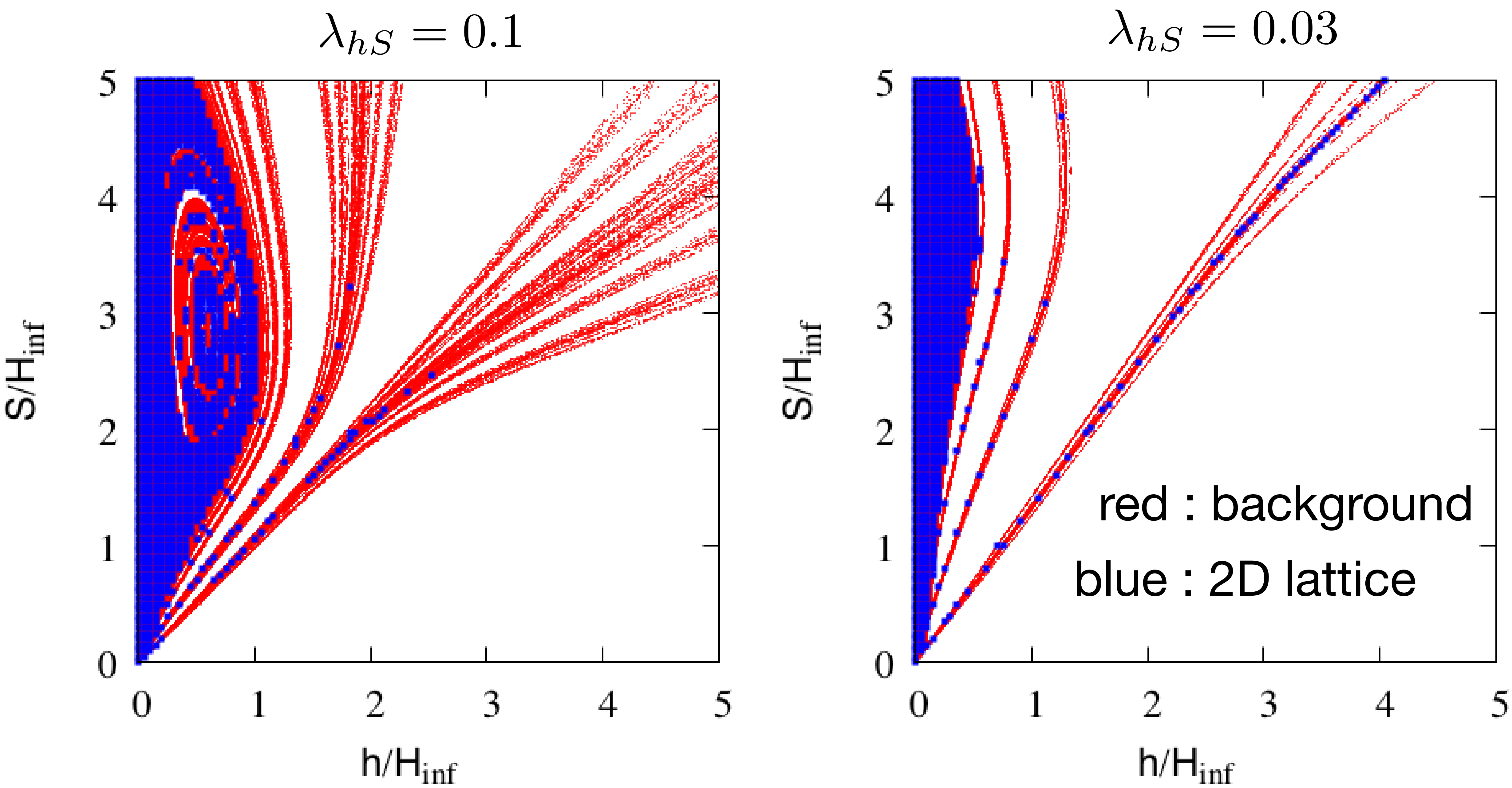
$$\ddot{S} + 3H\dot{S} - \frac{\nabla^2 S}{a^2} + \lambda_S S^3 + \lambda_{hS} h^2 S = 0$$

initial condition : $h_i = h_{\rm inf}, \quad S_i = S_{\rm inf}$

Background evolution



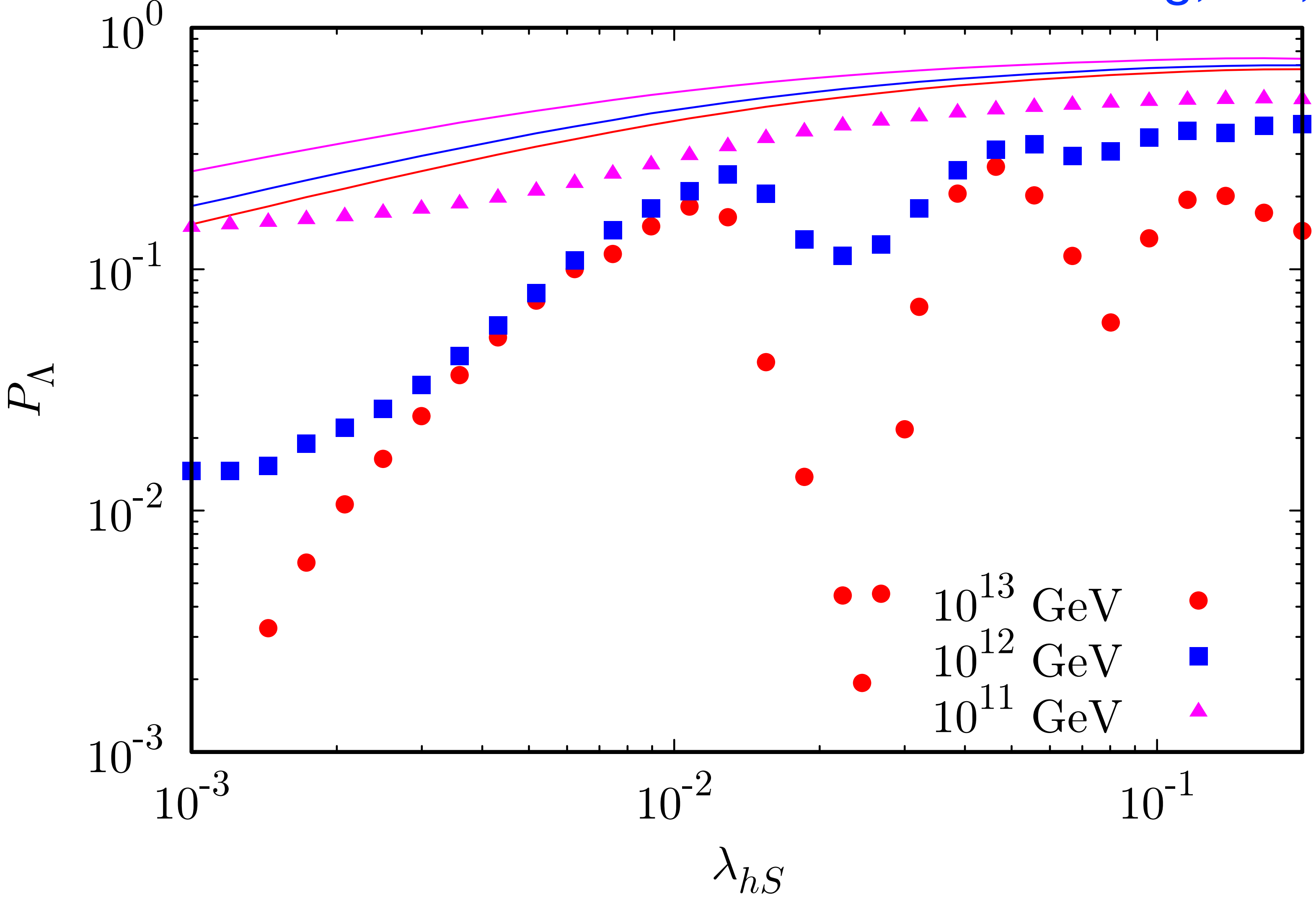
stability/instability chart



$$H_{\rm inf} = 10^{12}\ \text{GeV},\ \lambda_S = 0.01$$

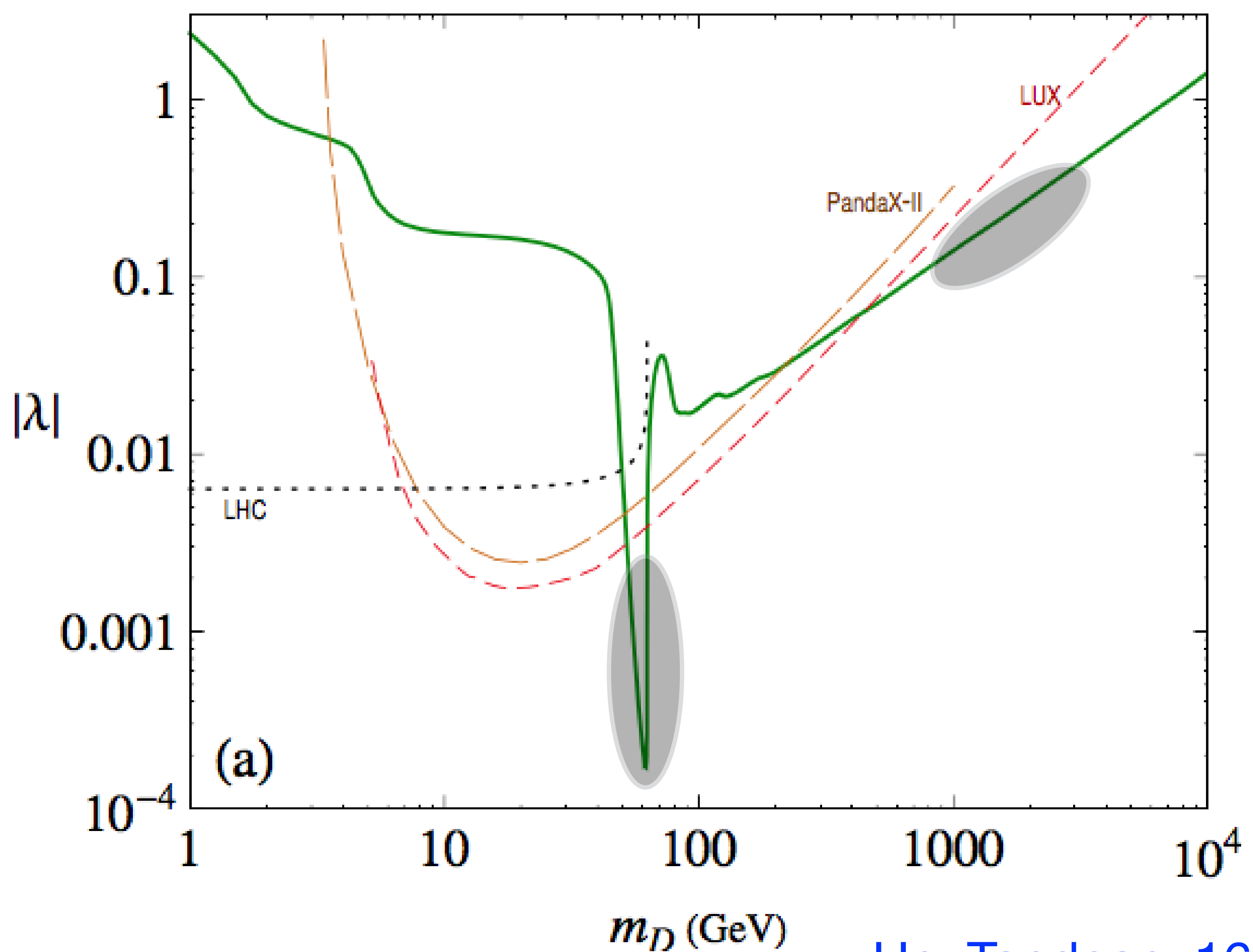
Survival probability after preheating

J.O. Gong, NK, 1705.11178



$$N = 60, \ \lambda_S = 0.01$$

Singlet spectator field as dark matter



He, Tandean, 1609.03551

Summary

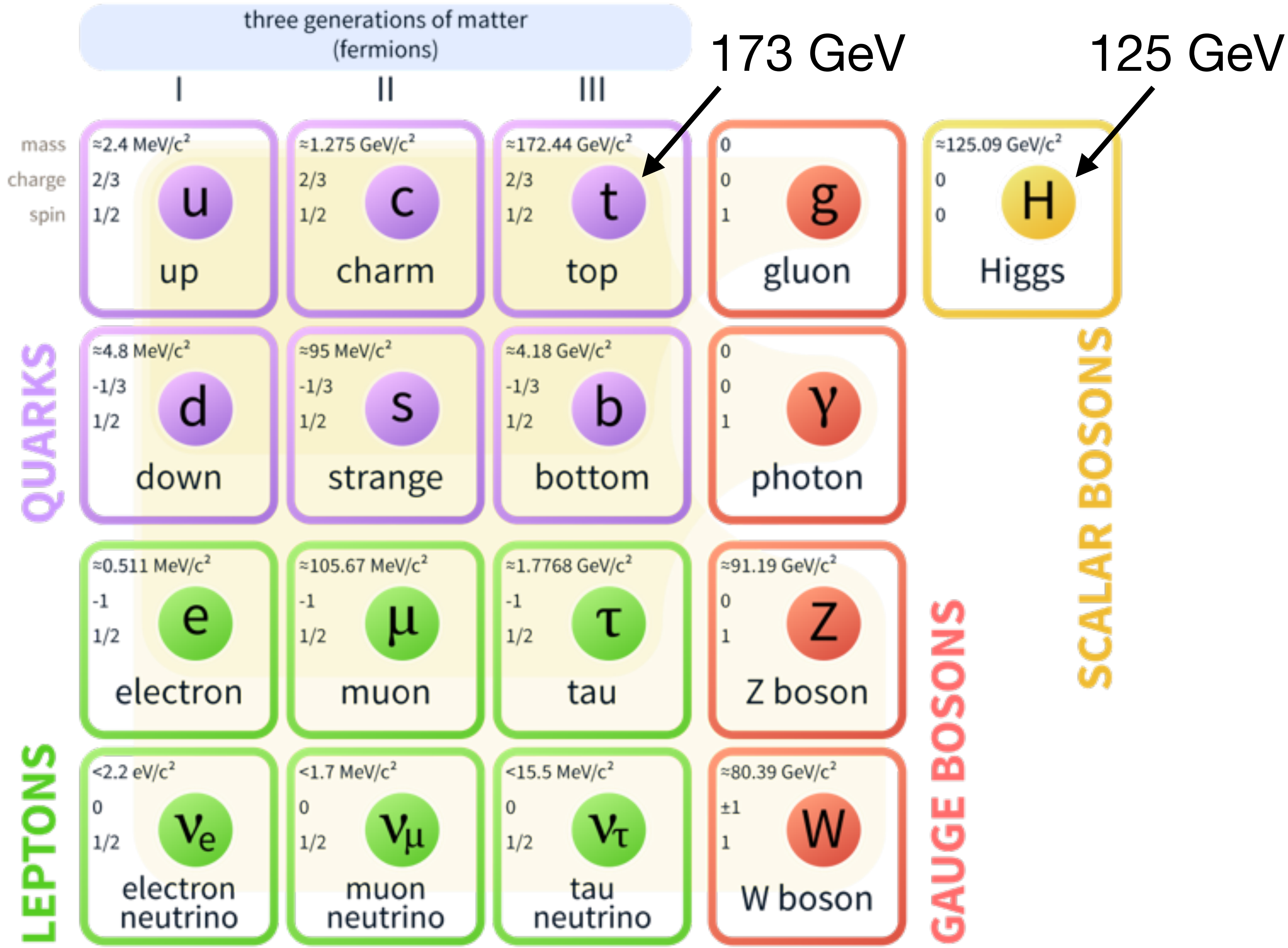
- We considered the singlet extension of SM
and solved multi-field stochastic dynamics during inflation
(Fokker-Planck equation)

- Stochastic dynamics of singlet (spectator) field
can largely enhance the survival probability during inflation

- Parametric resonance can occur in our model
but the resonance can be suppressed in a wide parameter range

- The singlet spectator field can be a dark matter
(marginally) consistent with the current constraint.

Standard Model of Elementary Particles



https://en.wikipedia.org/wiki/Standard_Model

RGEs (up to one loop order)

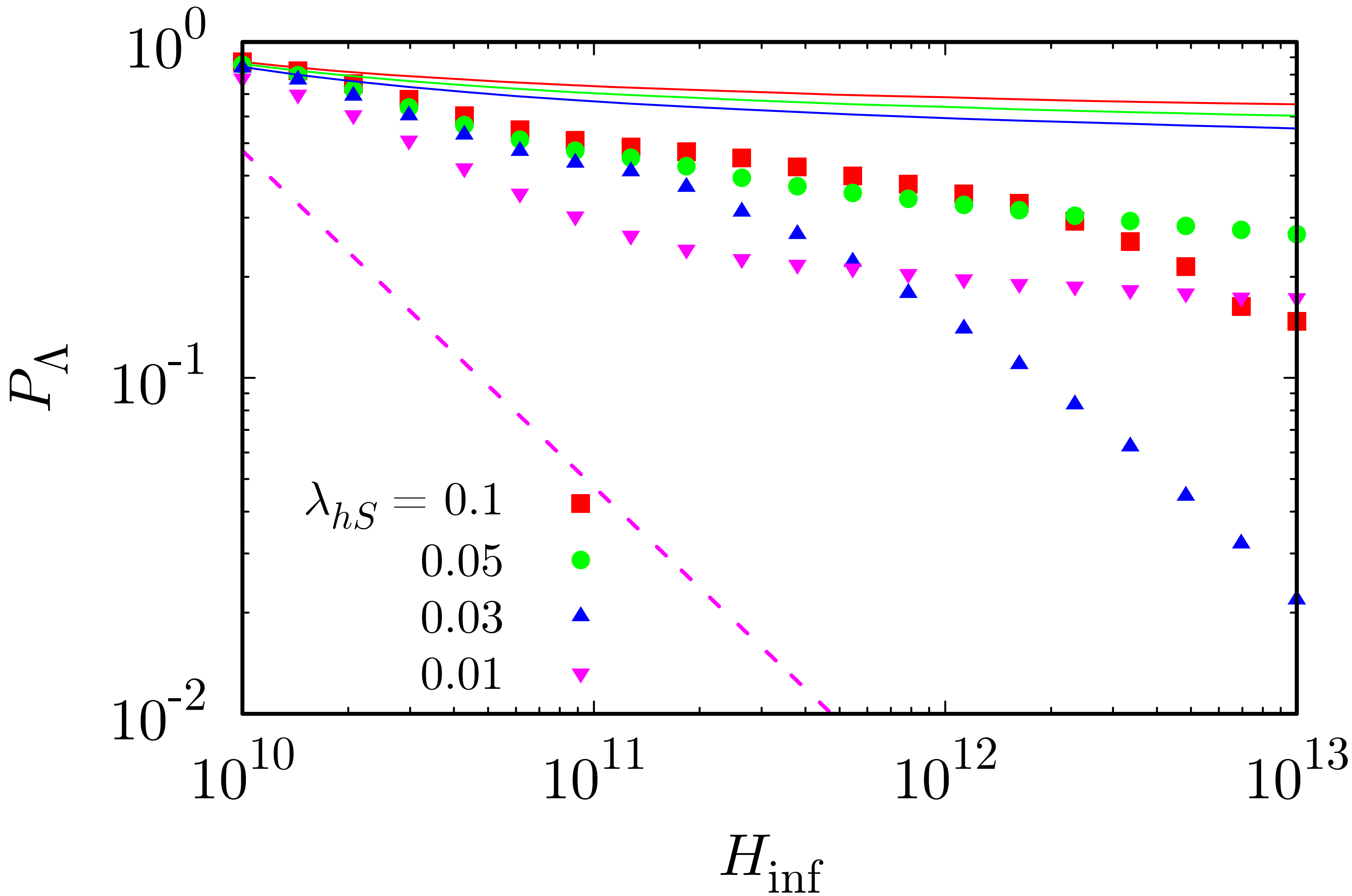
$$(4\pi)^2 \frac{d\lambda_h}{dt} = 24\lambda_h^2 + 12\lambda y_t^2 - 6y_t^4 - 3\lambda_h(g'^2 + 3g^2) + \frac{3}{8}[2g^4 + (g'^2 + g^2)^2] + \frac{\lambda_{hS}^2}{2}$$

$$(4\pi)^2 \frac{d\lambda_{hS}}{dt} = \lambda_{hS}\left[4\lambda_{hS} + 12\lambda_h + \lambda_S + 6y_t^2 - \frac{3}{2}(g'^2 + 3g^2)\right]$$

$$(4\pi)^2 \frac{d\lambda_S}{dt} = 3\lambda_S + 12\lambda_{hS}$$

with $t = \ln(\mu/M_t)$

Survival probability after preheating



$N = 60,\ \lambda_S = 0.01$